

Fall Meeting Call for Papers Update 614 – 615

J. A. Joselyn, J. Hirman, and G. R. Heckman

In the October 9, 1979 issue of *Eos*, Tsurutani and Baker (1979) expressed the tangible benefits of obtaining selected interplanetary data from the ISEE 3 spacecraft in real time. As shown in Figure 1, reproduced from Tsurutani and Baker, ISEE 3 is positioned between the earth and the sun in a halo orbit about the sun-Earth libration point. This location is well suited to provide advanced warning of the onset of geomagnetic disturbances. Although the data were not originally intended for real-time use, NOAA and NASA began to cooperate early in 1979 to solve the technical and administrative problems requisite to acquiring the raw data before editing and transmission to the experimenters. In March 1980 the data stream began arriving at the Space Environment Services Center (SESC), Boulder, Colorado. Daily data coverage varies but averages near 80%. The data are now being routinely used to support SESC military and civilian customers and the scientific community at large.

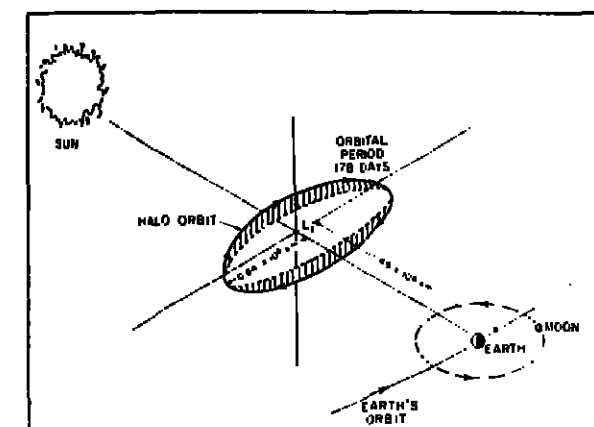


Fig. 1. The International Sun Earth Explorer 3 orbit about the sun-Earth libration point L_1 .

ISEC (a joint operation of NOAA and the Air Force Air Weather Service) is an around-the-clock center that closely monitors solar and geophysical conditions, issues warnings and alerts of special events, and writes twice-daily summaries and forecasts of the overall level of activity. Real-time and near-real-time data are routinely obtained from a chain of cooperating solar optical and radio observatories, including the remaining stations of the International Magnetospheric Study Program. Although interplanetary data have been provided in the past by Pioneers 6-9 and Vela spacecraft, and are now available by request from the principal investigators of Pioneers 6-9 and the Pioneer Venus Orbiter (depending upon a favorable location of Venus), ISEC 3 provides a vital platform for measuring ambient conditions in the solar wind some tens of minutes before Earth is impacted. In the past few months, this capability was especially helpful in providing support to NASA for the first flight of the shuttle *Columbia* and to the international European Energy Budget Campaign organized by D. Offer-

synoptic information. The general azimuthal direction of the magnetic field (e.g., toward or away from the sun) and the velocity are used in comparison with solar disc magnetograms from the Kitt Peak National Observatory to determine the large-scale source of the solar wind. Sector boundaries and other discontinuities in the direction of the solar wind are especially obvious. High-speed, low-density solar wind streams can be identified with specific coronal holes observed in helium 10830-Å spectroheliograms transmitted daily from Kitt Peak.

At an event-mode data peak, the ISEE 3 solar wind data can identify abrupt interfaces and shock waves in the solar wind. When these shocks traverse the magnetosphere, they are seen at geosynchronous satellite and low-latitude ground-based geomagnetic observatories as sudden impulses in the horizontal component, which may be storm sudden commencements. Figure 2 illustrates a sequence of observations on July 17, 1980. The discontinuity in the total interplanetary magnetic field at ISEE 3 occurred at 1840 UT (solar wind density and velocity data were not available). Previous to the detection of this shock, we had noted enhanced noise in the 3-kHz plasma wave experiment. Those emissions are due to instabilities driven by energetic protons traveling upstream of the shock and are often seen to precede intervals of shocked or highly disturbed conditions in the solar wind [Scarf, 1977; Kennel *et al.*, 1981]. Fifty minutes later, at 1930 UT, the impulse was registered at the low-latitude IMS stations. The delay time for the distance from ISEE to Earth of 1.5×10^6 km implied an assumed constant shock velocity of 500 km/s and an extrapolated delay from the sun to Earth of 3.5 days. This travel time corresponds with the occurrence of an X1/1 bright flare on the sun on July 14 at 0830 UT. From March 21, 1980, through April 30, 1981, 30 apparent shocks at ISEE 3 have been identified with sudden impulses on the ground. Twenty-three impulses in the geomagnetic field occurred at times when ISEE 3 data were not available. However, not all shock signatures seen at ISEE 3 can be identified with magnetic impulses at Earth, and not all sudden impulses on the ground can be identified in ISEE 3 data. Many shocks cannot be readily associated with a specific flare or other solar event, such as a filament disappearance. Further, the details of the shock at ISEE 3 do not seem to correspond with the details of the impulse at the ground. This variety in the event data exemplifies the complexity of the solar wind and the interaction between the solar wind and the magnetosphere.

Finally, the ISEE 3 solar wind data are a potentially quantitative predictor of geomagnetic storms and substorms. The key to quantitative prediction is an understanding of the mechanism of energy coupling between the solar wind and the magnetosphere. Numerous algorithms relating solar wind parameters to geomagnetic indexes have been proposed. (For discussions, see Crooker [1975]; Russell [1980]; and Donnelly [1979].) We have presently implemented only two of these predictors. The first algorithm is Amaldi's [1971] integration of the southward component of the interplanetary field. As explained above, this value is

Data Channel	Energy Range
<i>ISEE-3</i>	
SC 1	12-20 keV
SC 2	20-36 keV
SC 3-4	36-52 keV
<i>NOAA/GOES</i>	
'long'	1.5-12 keV
'short'	3-24 keV

calculated as a simple summation of the magnitude of the southward field and is displayed in units of gamma minutes in real time. The integration is terminated during those times when ISEE data are not received or when B_z turns northward, restarting when the data resume. Although no permanent record of this parameter has been kept, it has proved to be of value as an indicator of the intensity of geomagnetic activity. Our experience is in accordance with the well-known result that predominately northward fields are associated with very little geomagnetic activity; fluctuating fields are associated with minor disturbances; consistently southward fields are associated with active conditions (K values at Boulder of 3, 4, and occasionally 5); and strong southward fields ($B_z > -10$ γ and $2B_z/t > 1000$ γ min) are associated with storm levels. After a disturbed period, if the field turns and stays northward, magnetic activity diminishes.

A second parameter presently calculated and displayed is 'epsilon', a function first suggested by *Perreault and Akasofu* [1978] that is proportional to solar wind velocity and the square of the total interplanetary field and is strongly weighted toward southward fields. Detailed studies of this parameter relate it to AE, which is a global measure of geomagnetic activity substorms in the northern auroral zone. AE is not available in real time, even as an estimate, since SEEC auroral zone observatories are concentrated in Alaska, with our easternmost data arriving as a summary report every 90 min from Upper Heyford, England. Nevertheless, epsilon performs as a reasonable estimator of geomagnetic activity in that values greater than the threshold of 10^{10} erg/s are associated with significant geomagnetic activity [Akasofu, 1980]. There are other functions of interplanetary parameters that have been suggested as predictor algorithms. Although most are intended to provide warnings on the order of tens of minutes (i.e., the travel time from ISEE 3 to Earth plus some additional lag time for substorms), some can offer predictions of up to hours. These include the 3-kHz plasma wave noise observations mentioned above and subtle density and velocity variations that may signal the approach of a stream interface region in the solar wind (*Gosling et al.* [1978]; R. L. Rosenberg, private communication, 1980).

In summary, SESC greatly appreciates the spirit of exploration and cooperation that led to the acquisition of the real-time interplanetary data from ISEE 3. We are using

Irradiance measurements have large error bars near the centers of the absorption bands. The results imply that the 184-200 nm solar irradiance which penetrates to the stratosphere can be computed to an accuracy of ± 30 percent or better using presently available cross-sections. (Solar Irradiance, absorption).

J. Geophys. Res., Green, Paper 1C1045

0499 General or miscellaneous
MEASUREMENT OF VOLT/METER VERTICAL ELECTRIC FIELD
IN THE MIDDLE ATMOSPHERE
R. C. Maynard (Laboratory for Extraterrestrial
Physics, Goddard Space Flight Center, Greenbelt,
MD 20771), C. L. Crosskey, J. D. Mitchell and L.
C. Hale (Ionosphere Research Laboratory, The
Pennsylvania State University, University Park, PA
16802).

The flight for a series of wide-area atmospheric electrodynamics rockets was launched from Wallops Island, Virginia, at 7:18 EST on May 10, 1969. The mother-daughter configuration contained a vertical asymmetric double probe electric field instrument mounted vertically on top of the daughter payload, and a gradient coilometer and a single-axis (vertical) asymmetric double probe electric field instrument mounted horizontally on the side of the daughter payload. The rockets were launched by a Thor-Delta C vehicle which reached an apogee of 111 km and data were gathered from all instruments on the downing.

A vertical electric field with a maximum amplitude about 100 kV/m was observed in a layer between about 57 and 67 km. The magnetic potential across this layer was approximately 20 mV/km. The rocket measurements indicated that from 57 to 67 km there were no significant large-scale electric fields; however, the decrease in current density was insufficient to maintain vertical current continuity through the layer. These results establish the existence of a narrow magnetospheric electric fields, supporting previous results from single-satellite measurements. (Geophysics Res. Lett., 1970)

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0900 Transmutance
THE PHYSICAL FOUNDATION OF FORMATION LITHOLOGY
LOGIC WITH GANGA RAYS
Re 3583 N rays, gamma rays, and cosmic rays
W. Barzantj (Department of Physics, Massachusetts
Institute of Technology, Cambridge, MA 02139) B. V.
Kilim, and J. S. Wahl
We develop a theory for the gamma-ray spectrum in
the scattering and absorbing medium. Representations
are derived for the spectrum when sources are uniformly
distributed in an infinite medium. We express the
view that the formation of the spectrum is a result
of the interaction of the gamma-ray with the Compton
degradation of high-energy photons which are trans-
ported from the source to the neighborhood of the
point of interest. This study is extended to the
theory of the spectrum in an infinite medium, as
well as to a geometry appropriate for well logging.
Confirming evidence was shown Monte Carlo simula-
tions of the proposed model. The theory has applica-
tion to a well logging device for measurement of gamma-ray
absorption via the photoelectric effect, a par-
ticularly sensitive lithology.
GEOPHYSICS, VOL. 50, NO. 10, 1975, P. 10

0999 General or miscellaneous
DELINIZATION OF A LOW-VELOCITY BODY UNDER THE
ROOSEVELT MOUNTAIN SPRINGS GEOTHERMAL AREA, UTAH,
TELESEISMIC P-WAVE DATA
BY
JUN 1980 Thru April 1981
BSE 8130 East Rio
Jurnal Robinson (Geophysics Research, DSR, Box
12345, Boulder, Colorado) M. K. Yur
To assess the nature of the heat source associ-
ated with the Roosevelt Mt. Springs low-velocity structure
of the crust and uppermost mantle in the vicinity
of the Mineral Mountain, south of the volcanic region of
the geologic and geothermal activity of the Roosevelt
activity. A roughly square (30 X 30 km) array of
stations, centered on the low-velocity structure, was
operated for 72 hours, during which 72 teleseismic
events were recorded with sufficient resolution to
determine the nature of the low-velocity structure.
The results of the analysis of the data are presented
relative residual, using the array average for each
event as reference, show a clear pattern of the
existence of a localized region of relatively
low-velocity material extending up to 10 km from the
Mineral Mountain. A three-dimensional (3-D) inversion
of the data confirms this conclusion, and shows that
the low-velocity material extends to a depth of 7 percent
less than the surrounding rock) centered under the
geothermal area and extending from about 10 km from
the Mineral Mountain to the south. The low-velocity
structure obtained in the inversion clearly reflects
the structure of the region, part of the basement
range province. An essentially changing thickness
of the low-velocity material is observed in the central
Mineral Mountains, indicates the presence of a
small but intensely anisotropic region of about 10
km from the Mineral Mountain. The low-velocity
structure is not a simple, uniform, and isotropic
structure found in the region of preference.
The low-velocity material is a partial melt
region in terms of approximately 10 percent partial
melting. The low-velocity material is a partial melt
region in terms of approximately 10 percent partial
melting. The low-velocity material is a partial melt
region in terms of approximately 10 percent partial
melting.

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